

New reaction rate for $^{16}\text{O}(p, \gamma)^{17}\text{F}$ and its influence on the oxygen isotopic ratios in massive AGB stars

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The $^{16}\text{O}(p, \gamma)^{17}\text{F}$ reaction rate is revisited with special emphasis on the stellar temperature range of $T = 60$ – 100 MK, important for hot bottom burning in asymptotic giant branch (AGB) stars. We evaluate existing cross-section data that were obtained since 1958 and, if appropriate, correct published data for systematic errors that were not noticed previously, including the effects of coincidence summing and updated effective stopping powers. The data are interpreted by using two different models of nuclear reactions, that is, a potential model and R-matrix theory. A new astrophysical S factor and recommended thermonuclear reaction rates are presented. As a result of our work, the $^{16}\text{O}(p, \gamma)^{17}\text{F}$ reaction has now the most precisely known rate involving any target nucleus in the mass $A \geq 12$ range, with reaction rate errors of about 7% over the entire temperature region of astrophysical interest ($T = 0.01$ – 2.5 GK). The impact of the present improved reaction rate with its significantly reduced uncertainties on the hot bottom burning in AGB stars is discussed. In contrast to earlier results we find now that there is not clear evidence to date for any stellar grain origin from massive AGB stars.

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I. INTRODUCTION

The $^{16}\text{O}(p, \gamma)^{17}\text{F}$ reaction is characterized by a number of exceptional attributes. At lower bombarding energies it provides a textbook example for a nonresonant reaction cross section because the lowest-lying resonance is located at a relatively high laboratory energy of 2.66 MeV [1]. The absence of low-energy resonances and the high binding energy of ^{16}O are among the main reasons for the fact that the cross section can be described in terms of simple nuclear reaction models. Indeed, the $^{16}\text{O}(p, \gamma)^{17}\text{F}$ reaction is a prime example for the direct capture reaction model that assumes that the projectile is captured via a single-step process into a final-state orbit outside a closed ^{16}O core (see, for example, Ref. [2]). The absence of low-energy resonances is also the reason for the fact that the $^{16}\text{O}(p, \gamma)^{17}\text{F}$ reaction is the slowest process among all the proton-induced reactions in the CNO target mass region [3]. This reaction has been studied many times at low energies. The experimental techniques applied include the activation method [4,5], the in-beam study of prompt γ rays [2,6,7], and measurements in inverse kinematics [8]. It is generally assumed that the different measurements are in agreement. Despite these facts, the thermonuclear reaction rates evaluated by the NACRE collaboration have relatively large errors amounting, for example, to +35% and -43% at $T = 0.06$ – 0.1 GK [9].

Many stars, including the Sun, will eventually pass through an evolutionary phase that is referred to as the asymptotic giant branch (AGB; see Ref. [10] for a review). This phase involves a hydrogen and a helium shell that burn alternately surrounding an inactive stellar core. Convection carries the products of

nucleosynthesis to the stellar surface where material is ejected via strong stellar winds [11]. A fraction of the ejected matter condenses in the form of small grains. Some of the grains traveled the interstellar medium and have been incorporated into primitive meteorites at the birth of the solar system. Such stellar grains can be extracted in the laboratory from their host meteorites and their isotopic composition can be analyzed with high precision [12]. The measured isotopic abundance ratios deviate substantially from solar system abundance ratios, which result from a homogeneous mixture of contributions from countless stars prior to solar system formation. Stellar grains, instead, exhibit isotopic abundance ratios that are characteristic for the composition of their parent stars, and hence they most likely provide strong constraints for stellar models. The present work has been motivated by studies showing that variations of the $^{16}\text{O}(p, \gamma)^{17}\text{F}$ reaction rate influence sensitively the $^{17}\text{O}/^{16}\text{O}$ isotopic ratio predicted by models of massive ($\geq 4M_{\odot}$) AGB stars, where proton captures occur at the base of the convective envelope (hot bottom burning, HBB). Specifically, a recent study demonstrated that a fine-tuning of the $^{16}\text{O}(p, \gamma)^{17}\text{F}$ reaction rate may account for the measured anomalous $^{17}\text{O}/^{16}\text{O}$ abundance ratio in the extraordinary presolar spinel grain OC2 [13].

In the present work we focus our attention on a detailed evaluation of existing $^{16}\text{O}(p, \gamma)^{17}\text{F}$ cross-section data (Sec. II). This level of detail was not practical in the broad NACRE compilation, which includes many different reactions. We interpret the evaluated data in terms of two different nuclear reaction models (Sec. III) and extract astrophysical S factors, reaction rates, and associated errors (Sec. IV). The new $^{16}\text{O}(p, \gamma)^{17}\text{F}$ thermonuclear reaction rates are then

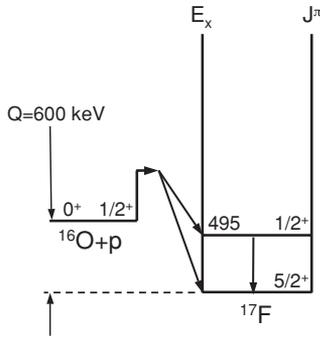


FIG. 1. Nuclear energy level diagram for $^{16}\text{O}(p, \gamma)^{17}\text{F}$. The level parameters and the reaction Q value are adopted from Refs. [1] and [14], respectively.

incorporated into stellar models to derive improved oxygen isotopic ratios from the hot bottom burning in AGB stars (Sec. V).

II. EVALUATION OF EXISTING DATA

The radiative proton capture on ^{16}O ($Q = 600.27 \pm 0.25$ keV [14]) can proceed via $E1$ radiation either to the ground state ($J^\pi = 5/2^+$) or to the first excited state ($E_x = 495.33 \pm 0.10$ keV; $J^\pi = 1/2^+$) of ^{17}F [1]. A nuclear level diagram is shown in Fig. 1. In this section, we evaluate the existing cross-section data and discuss certain corrections that we deem necessary. Notice that we make a distinction between experiments that measure separately the transitions to the two final states and those that report only on the sum contribution of both transitions. The data will be displayed in terms of the astrophysical S factor, defined as

$$S(E) = E\sigma(E)e^{2\pi\eta}, \quad (1)$$

with σ the reaction cross section, $e^{-2\pi\eta}$ the Gamow factor, and η the Sommerfeld parameter. Throughout this experimental section, all kinematic quantities are given in the laboratory system, unless mentioned otherwise.

A. Data of Chow, Griffiths, and Hall (1975)

The work of Chow, Griffiths, and Hall [6] reports on cross sections for $^{16}\text{O}(p, \gamma)^{17}\text{F}$ at seven bombarding energies in the range of $E_p = 845\text{--}2554$ keV. The cross sections are separately given in their Table 4 for the transitions to the ground state (four data points) and first excited state (seven data points) of ^{17}F . It is important to emphasize that these authors normalized the radiative capture cross section to the Rutherford scattering cross section at $E_p = 405$ keV. This procedure provides capture cross sections independently from the target stoichiometry, stopping powers, and beam straggling and thus avoids many normalization errors that are frequently associated with absolute cross-section determinations. For more information, see Ref. [3].

Not all the experimental details are provided in Ref. [6] to assess the quoted errors. Some of their reported cross-section errors are as small as $\pm 3\%$. However, the elastic-scattering measurement alone contributes an error of about $\pm 2\%$, as can be seen from their Table 1. Based on the geometry of the

experiment and the methods applied, the following estimates for additional error contributions seem reasonable: $\pm 3\%$ for the γ -ray efficiency; $\pm 1\%$ for the escape peak detection; $\pm 1\%$ for the angular distribution correction; and $\pm 3\%$ for the determination of the effective bombarding energy in their 25-keV-thick solid oxygen targets. Other sources of errors, such as corrections for the finite solid angle of the γ -ray detector, are mentioned in Ref. [6]. The quadratic addition of the above-mentioned errors gives a value of $\pm 5\%$ and, therefore, we suspect that the smallest of the capture cross-section errors quoted in their Table 4 have been somewhat underestimated. As a result of the above discussion, we adopt errors of at least $\pm 5\%$ for their quoted capture cross sections. The modified data of Chow, Griffiths, and Hall [6] are displayed in Fig. 2.

B. Data of Morlock *et al.* (1997)

The measurement of Morlock *et al.* [7] (see also Ref. [15]) was performed in the energy range of $E_p = 280\text{--}3740$ keV by using an extended gas target. Their $^{16}\text{O}(p, \gamma)^{17}\text{F}$ cross

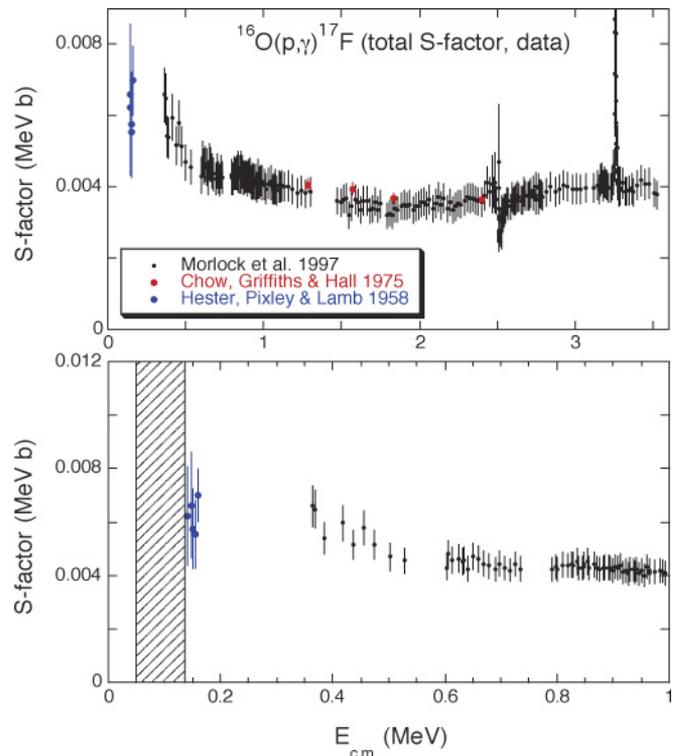


FIG. 2. (Color online) Total astrophysical S factor for the $^{16}\text{O}(p, \gamma)^{17}\text{F}$ reaction versus center-of-mass bombarding energy; (top) all data considered and (bottom) expanded view for low energies. A number of different corrections are applied to the data of Refs. [4,6,7], whereas the results of Refs. [2,5] are disregarded in the present work; see text for details. Note that we neglected the lowest energy data point of Ref. [4] because the associated error is very large ($\pm 52\%$). The total S factor varies slowly, except near the vicinity of the lowest-lying resonances at $E_p = 2.66$ and 3.47 MeV [1]. The shaded area in the lower part indicates the effective energy range of stellar burning in massive AGB stars corresponding to temperatures of $T = 60\text{--}100$ MK.

section was determined relative to the $^{16}\text{O}(p, p)^{16}\text{O}$ elastic-scattering cross section, which, in turn, was normalized to the Rutherford scattering of protons on heavy noble gases (krypton and xenon). A number of effects were considered by Refs. [7,15] in the analysis, including the energy loss and straggling of protons in the gas target, the energy spread of the beam, the finite solid angle of the γ -ray detector, the angular distribution of the emitted γ rays, and the dependence of the detection efficiency on γ -ray energy and emission angle.

We noticed, however, that the data of Refs. [7,15] were not corrected for coincidence summing. This effect arises because the primary γ ray populating the $E_x = 495$ keV level and the secondary γ -ray from the subsequent de-excitation of this state are in time coincidence (Fig. 1). Because both photons are emitted nearly simultaneously, the detector is not able to distinguish between the two but instead records a single event with the summed energy. This experimental artifact gives rise to too many counts for the ground-state transition (“summing-in”) and too few counts for the transition to the first excited ^{17}F state (“summing-out”). The effect can be corrected for only with the proper knowledge of both the peak and total detection efficiencies (see Ref. [3] for details). Most of these quantities have not been reported in Refs. [7,15] but can be deduced from other sources, as described below.

We adopted the peak efficiencies from Ref. [15] and from Refs. [16,17]. The latter studies were performed with a similar detection setup as in Morlock *et al.* [7]. To estimate the total detection efficiency, we performed GEANT simulations for a similar setup involving a germanium γ -ray detector of 100% relative efficiency. We find summing-out correction factors of 1.06 ± 0.02 for the primary transition to the first excited ^{17}F state and 1.05 ± 0.02 for the secondary de-excitation of the 495-keV level. For the primary ground-state transition, the summing-in depends on the cross section of the transition $\text{DC} \rightarrow 495$. The correction reduces the reported cross section at the lowest measured energies (few 100 keV) by about 10%, whereas the correction is negligible at the highest energies (several MeV).

We found evidence in Morlock’s thesis [15] that an additional error of 10% has to be applied to all data points. Note that this additional error was disregarded in Ref. [7]. We also discovered an inaccuracy (too few significant digits) in the numerical conversion from cross section to S factor presented in Ref. [15]. The modified data of Morlock *et al.* [7] and Morlock [15] are displayed in Fig. 2.

C. Data of Hester, Pixley, and Lamb (1958)

Hester, Pixley, and Lamb [4] measured the total $^{16}\text{O}(p, \gamma)^{17}\text{F}$ cross section at six energies in the bombarding energy range of $E_p = 140\text{--}170$ keV. They measured the ^{17}F activity after exposing an infinitely thick oxygen target to an intense proton beam. Thus they could not distinguish between transitions to the ground or first excited state in ^{17}F . Their study represents an absolute cross-section measurement, in the sense that the deduced cross section depends on absolute detection efficiencies, absolute incident charge integration, stopping powers, and target stoichiometry. Recall that many

of the systematic errors related to these sources are avoided in the studies of Refs. [6,7] mentioned above.

The cross section of Ref. [4] was obtained from the expression

$$\sigma(E) = \frac{Z_t \epsilon_{\text{eff}} Y}{2E^{3/2}} \left(1 + \frac{\sqrt{E}}{Z_t} + \dots \right) \quad (2)$$

with Z_t the charge of the target nucleus, E the laboratory bombarding energy in MeV, Y the thick target yield, and ϵ_{eff} the effective stopping power. This expression is derived under the following assumptions: (i) the target is infinitely thick; (ii) the S factor is approximately constant in the energy range of the measurement; and (iii) the stopping power is constant over the effective thickness of the target. For the effective stopping power of their Al_2O_3 target, $\epsilon_{\text{eff}} = \epsilon_{\text{O}} + (2/3)\epsilon_{\text{Al}}$, they used a value of $\epsilon_{\text{eff}} = 2.8 \times 10^{-14}$ eV cm^2 that was adopted from an unpublished report [18]. We corrected the measured cross sections of Ref. [4] by using modern stopping powers derived from the code SRIM [19]. The correction decreases the cross section reported in Ref. [4] by 8% (2%) at the highest (lowest) energy measured. The modified data of Hester, Pixley, and Lamb [4] are displayed in Fig. 2.

D. Data of Becker *et al.* (1982)

Becker *et al.* [8] measured the cross section for radiative capture on ^{16}O at a center-of-mass (c.m.) bombarding energy of $E_p^{\text{c.m.}} = 853$ keV. Their study was performed in inverse kinematics by directing a ^{16}O heavy-ion beam onto an extended windowless hydrogen gas target. It must be emphasized that Becker *et al.* performed similar measurements for other nuclear reactions (see Table 1 in Ref. [8]) and that their reported resonance strengths are frequently used in the literature as absolute strength standards (see, for example, Refs. [3,20,21]). For the $^{16}\text{O}(p, \gamma)^{17}\text{F}$ cross section at $E_p^{\text{c.m.}} = 853$ keV they report a value of $\sigma = (0.92 \pm 0.13) \times 10^{-3}$ mb. Although not clearly stated in their article, this value does not correspond to the total cross section but represents the cross section for the transition to the first excited ^{17}F state only [22]. As will be shown later, the data point of Becker *et al.* [8] agrees with the data of Chow, Griffiths, and Hall [6] and Morlock *et al.* [7].

E. Other data

Two more studies [2,5] report low-energy cross sections for the $^{16}\text{O}(p, \gamma)^{17}\text{F}$ reaction. These report total cross sections and provide no information on individual transitions to the two final ^{17}F states. We disregarded these data for the reasons given below.

The study by Tanner [5] reports cross sections in the bombarding energy range of $E_p = 274\text{--}616$ keV. The absolute cross-section measurement was performed with a WO_3 target by using the activation method. At the time, stopping powers for tungsten were not available and, as the closest approximation, tabulated values for tantalum were used instead [23]. Clearly, these stopping powers are several decades old and the

cross sections reported by Tanner [5] should in principle be corrected by using modern stopping power values for oxygen and tungsten. However, our attempt at a correction was futile, mainly because Tanner obtained the cross section from the yield and the effective stopping powers of his 150-keV-thick target by numerical integration. Because the measured yields are not provided in Ref. [5] there is no obvious way to correct his reported cross sections. Consequently, we are compelled to disregard the data of Tanner [5] in our analysis.

Rolfs [2] reports total $^{16}\text{O}(p, \gamma)^{17}\text{F}$ cross sections in the bombarding energy range of $E_p = 0.3\text{--}3.0$ MeV. He did not determine the absolute scale of the cross section but normalized his results relative to the absolute cross section of Tanner [5] that was measured at an energy of $E_p = 616$ keV. Because we did not succeed in correcting the latter data for improved stopping powers (see above), we are also compelled to disregard the data of Rolfs [2] in our analysis. Nevertheless, the data of Ref. [2], which were measured over a broad energy range, demonstrated the rise of the S factor at low energies and the measured energy dependence agrees with later results [7].

III. NUCLEAR REACTION MODEL CALCULATIONS

In this section we describe the fitting of a number of models to the data reported in Chow, Griffiths, and Hall [6] and Morlock *et al.* [7]. The fits are performed for each transition separately. No attempt has been made in the present work to constrain the fitting by using the reported total S factors of Hester, Pixley, and Lamb [4]. However, the latter data, together with the single data point of Becker *et al.* [8] for the transition to the first excited ^{17}F state, provide important cross checks for the present results. One of our main goals is to extrapolate the S factor to center-of-mass energies of $E \approx 50\text{--}140$ keV, corresponding to stellar temperatures of $T \approx 0.06\text{--}0.1$ GK that are characteristic of hot bottom burning in AGB stars (see the shaded area in Fig. 2). It will be seen that our procedures give rise to significantly improved thermonuclear reaction rates over the entire temperature range of astrophysical interest ($T = 0.01\text{--}10$ GK). Throughout this theoretical section and the following sections, all kinematic quantities are given in the center-of-mass system, unless mentioned otherwise.

A proper analysis should distinguish between statistical and systematic errors. We corrected the data of Refs. [6,7] for a number of different effects (Sec. II). However, the original articles do not provide sufficient information to quantitatively disentangle statistical and systematic errors. In view of this ambiguity, we decided to treat the combined uncertainties as statistical errors in the fitting procedure. The same problem and strategy is encountered in Ref. [24]. More information regarding our error analysis, and a justification for the adopted procedure, is given in Sec. II.

The radiative capture cross section for $^{16}\text{O}(p, \gamma)^{17}\text{F}$ from an initial state $J_i\pi_i$ to a final state $J_f\pi_f$ is given by

$$\sigma_c(J_i\pi_i \rightarrow J_f\pi_f) = \frac{\pi}{2k^2}(2J_i + 1)|U^\gamma(J_i\pi_i \rightarrow J_f\pi_f)|^2, \quad (3)$$

where U^γ is the transition matrix element. For a given final state J_f , the total cross section is obtained by summing over all initial angular momenta J_i . This definition of the

cross section is common to all models; the only difference is the way of calculating the transition matrix element. In the potential model, this term is computed from an integral involving the initial and final wave functions, and the radial part of the electromagnetic operator. In the R -matrix approach, it is split in two parts, involving different parameters fitted to experiment. We give more detail below on both models.

A. Potential model

We start the analysis of existing data by using a simple model for the low-energy cross section in $^{16}\text{O}(p, \gamma)^{17}\text{F}$. The model is referred to as “direct proton capture” and it assumes a single-step process where the proton is directly captured, without formation of a compound nucleus, into a final bound state with the emission of a photon. Many nonresonant reactions in nuclear astrophysics have been described by this model (see Ref. [25] and Table I in Ref. [26]).

Numerous studies (see, for example, Ref. [26]) have shown that at relatively low bombarding energies and small binding energies of the final states the direct capture of protons occurs mainly far outside the nuclear radius. For this reason, the direct capture process is sometimes referred to as “extranuclear” capture. This also implies that the calculated cross section is expected to be relatively insensitive to the details of the model used to describe the nuclear interior.

In the potential-model formalism the transition matrix element is given by

$$U^\gamma(J_i\pi_i \rightarrow J_f\pi_f) \sim \sqrt{C^2\mathcal{S}} \int_0^\infty u_i(r, E) \mathcal{O}_{\omega\lambda}(r) u_f(r) r^2 dr, \quad (4)$$

where $C^2\mathcal{S}$ denotes the spectroscopic factor of the final state [27], and $\mathcal{O}_{\omega\lambda}$ is the radial part of the multipole operator for electromagnetic radiation of character ω (either electric or magnetic) and multipolarity λ , u_i denotes the radial part of the partial waves of the initial scattering state, and u_f is the radial wave function of final bound state. The model adopted here is similar to the one employed in Ref. [7]. The full formalism is given in Refs. [28,29].

For example, for the $^{16}\text{O}(p, \gamma)^{17}\text{F}$ reaction the transition to the ground state proceeds predominantly via $E1$ radiation and angular momenta of $\ell_i = 1, 3 (J_i = 3/2^-, 5/2^-, 7/2^-) \rightarrow \ell_f = 2 (J_f = 5/2^+)$ ($\text{DC} \rightarrow 0$), whereas the transition to the first excited state at $E_x = 495$ keV proceeds via $E1$ radiation and angular momenta of $\ell_i = 1 (J_i = 1/2^-, 3/2^-) \rightarrow \ell_f = 0 (J_f = 1/2^+)$ ($\text{DC} \rightarrow 495$). Other transitions, that is, those of $M1$ or $E2$ character, are negligible for the direct proton capture on ^{16}O . See also Ref. [30].

The scattering- and bound-state wave functions are generated by an optical potential

$$V(r) = \xi V_F(r) + \xi_{\text{s.o.}} \frac{4}{r} \frac{dV_F}{dr} \vec{L} \cdot \vec{S} + V_C(r), \quad (5)$$

where V_F denotes a folding potential and ξ is the folding potential strength parameter. Folding potentials have the major advantage that the geometry of the potential is fixed by the folding calculation. In other words, the potential can be changed only by a variation of the strength parameter ξ . The

(weak) spin-orbit potential term has the usual Thomas form and is characterized by the spin-orbit strength parameter $\xi_{\text{s.o.}}$, whereas the Coulomb potential V_C is given by a uniformly charged sphere of radius R_C . The folding potential is determined using the effective nucleon-nucleon interaction adopted from the well-established DDM3Y parametrization. For the volume integral per interacting nucleon pair and the root-mean-square radius we find values of $J_R = 525.93 \text{ MeV fm}^3$ and $r_{F,\text{rms}} = 3.311 \text{ fm}$, respectively, whereas we also adopt the latter value for the Coulomb radius. For details regarding the folding procedure, see Refs. [7,31].

For generating the bound-state wave function, the parameter ξ is adjusted to reproduce the binding energies of the ground and first excited state, whereas the parameter $\xi_{\text{s.o.}}$ can be constrained by the energy splitting of the lowest $5/2^+$ and $3/2^+$ states in ^{17}F . We find values of $\xi = 1.0976$ and $\xi_{\text{s.o.}} = -0.1757 \text{ fm}^2$. For generating the scattering-state wave function, it was shown in Ref. [7] that good agreement between the experimental capture data and the calculation is found by using a value of $\xi \approx 1.0$ and the same value for $\xi_{\text{s.o.}}$ as for the bound-state calculation. This potential also describes reliably the $^{16}\text{O}(p, p)^{16}\text{O}$ elastic-scattering data at low energies.

Once the folding potential strength parameter ξ for generating the scattering wave function has been fixed, the spectroscopic factor C^2S is the only adjustable parameter in the above model. Our strategy will be as follows. Initially, the potential parameter is held constant at $\xi = 1.0$ and the data of Chow, Griffiths, and Hall [6] and Morlock *et al.* [7] are fitted independently. The least-squares fits provide for each transition and data set the corresponding value of C^2S with an associated error. Afterward, the above procedure is repeated by systematically varying the value of ξ to investigate the sensitivity of the extrapolated S factor.

Numerous studies have shown that the spectroscopic factors for the transitions to the ground and first excited states in $^{16}\text{O}(p, \gamma)^{17}\text{F}$ are close to unity (see, for example, Ref. [26]). It must be emphasized, however, that we are here not at all concerned with the extraction of accurate spectroscopic factors from the measured low-energy cross sections. The magnitude of C^2S will depend on the potential parameter ξ , as is apparent from Eqs. (4) and (5). In the present analysis, the spectroscopic factors are simply regarded as intermediate results, or scaling factors, whose derived uncertainty determines partially the error in the extrapolated astrophysical S factor.

Our potential model does not account for the presence of resonances. Therefore, we consider only the center-of-mass energy range below $E = 2.4 \text{ MeV}$ for the least-squares fits described in this section. The fits to the experimental S -factor data for the transition to the ^{17}F ground state are shown in Fig. 3. The results are $C^2S(\text{DC} \rightarrow 0) = 1.09 \pm 4.0\%$ ($\chi_{\text{red}}^2 = 1.4$) for the data of Chow, Griffiths, and Hall [6] and $C^2S(\text{DC} \rightarrow 0) = 1.22 \pm 0.9\%$ ($\chi_{\text{red}}^2 = 1.1$) for the data of Morlock *et al.* [7]. The absolute normalization of these two data sets differs thus by 12%. Figure 4 shows the fits to the experimental S -factor data of the transition to the first excited state in ^{17}F . In this case, we find $C^2S(\text{DC} \rightarrow 495) = 1.05 \pm 1.7\%$ ($\chi_{\text{red}}^2 = 1.5$) for the data of Ref. [6] and $C^2S(\text{DC} \rightarrow 495) = 1.02 \pm 0.5\%$ ($\chi_{\text{red}}^2 = 0.3$) for the data of Ref. [7]. The difference in the absolute normalization of the

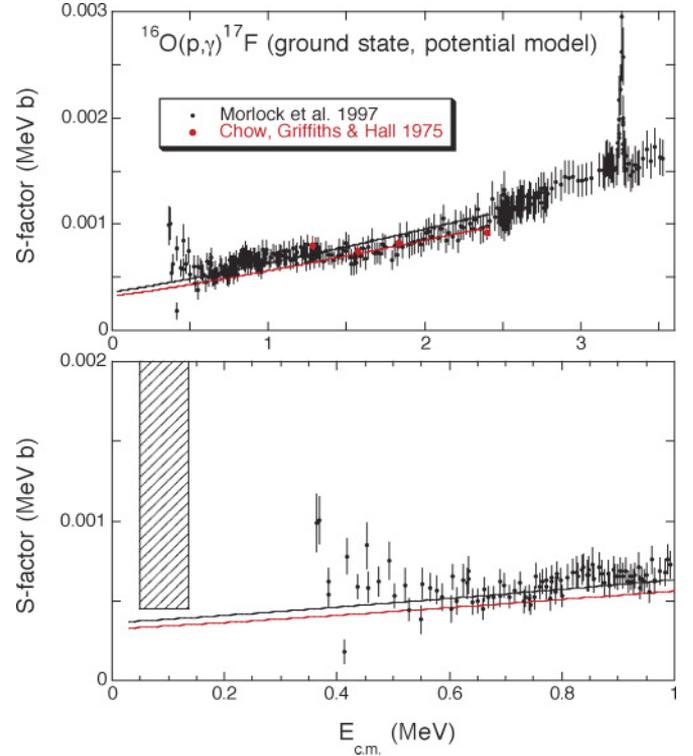


FIG. 3. (Color online) S factor for the proton capture into the ^{17}F ground state. The black and red lines represent fits to the data of Morlock *et al.* [7] and Chow, Griffiths, and Hall [6], respectively. The lines are obtained with a folding potential model (Sec. III A). The 68% confidence bands are rather narrow and are not displayed for reasons of clarity. Only data below a center-of-mass energy of $E = 2.4 \text{ MeV}$ are included in the fits. The shaded area in the lower part indicates the effective energy range for the hot bottom burning in AGB stars.

two data sets amounts to 3%. Note that the upturn in the S factor at low energies is well understood. See, for example, Refs. [6,7].

A different choice for the parameter ξ of the scattering potential changes not only the magnitude but also the energy dependence of the calculated direct capture S factor. The sensitivity of the extrapolated S factor to this parameter will be explored at a center-of-mass energy of $E = 90 \text{ keV}$, which is located at the center of the energy region important for the hot bottom burning in AGB stars (shaded areas in Figs. 3 and 4). We quote in the following the results obtained for the dominant transition to the first excited state in ^{17}F ($\text{DC} \rightarrow 495$). For this sensitivity study we use the corrected experimental data of Ref. [7] with their statistical errors only (that is, excluding the additional 10% overall uncertainty; see Sec. II B). Similar results are obtained for the data of Ref. [6]. For the reference calculation (using $\xi = 1$) we find a spectroscopic factor of $C^2S = 1.02$ and an S factor of $S_{\text{ref}}(90 \text{ keV}) = 7.03 \text{ keV b}$ for the $\text{DC} \rightarrow 495$ transition. For example, a variation in the potential strength parameter ξ by 10% changes both the spectroscopic factor obtained from a least-squares fit of the data of Morlock *et al.* and the extrapolated S factor at $E = 90 \text{ keV}$ by about 5%. The results of our sensitivity

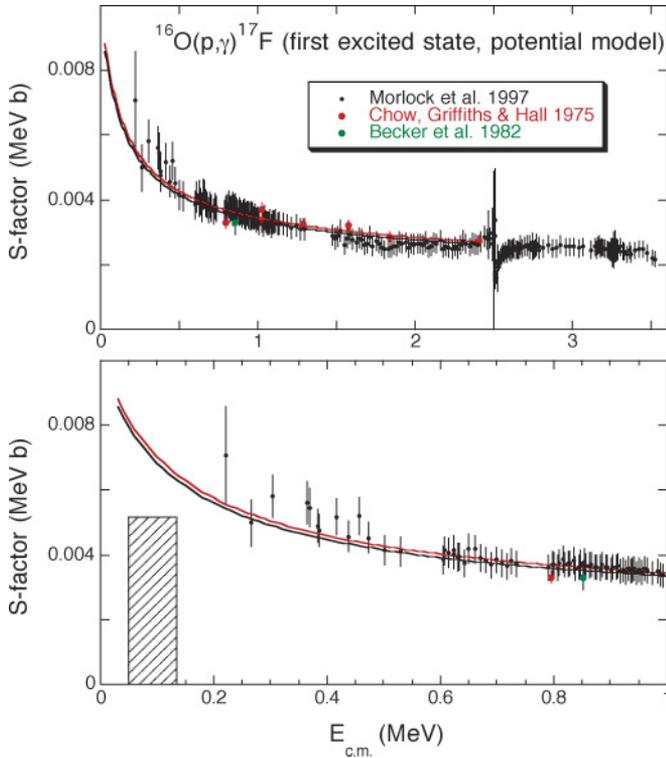


FIG. 4. (Color online) S factor for the proton capture into the first excited state of ^{17}F . The black and red lines represent fits to the data of Morlock *et al.* [7] and Chow, Griffiths, and Hall [6], respectively. The lines are obtained with a folding potential model (Sec. III A). Only data below a center-of-mass energy of $E = 2.4$ MeV are included in the fits. Notice the single data point reported by Becker *et al.* [8], shown in green. See also caption of Fig. 3.

study for center-of-mass energies of $E = 90$, 500, and 1000 keV are given in Table I. We also find that calculations with Woods-Saxon scattering potentials give similar results, even when the Woods-Saxon parameters are varied over a relatively broad range.

It is interesting to note that a variation in the potential strength parameter ξ influences the calculated energy de-

TABLE I. Astrophysical S factors for $^{16}\text{O}(p, \gamma)^{17}\text{F}$ at energies of $E = 90$, 500, and 1000 keV. The results shown here are obtained by fitting the data of Ref. [7] for the dominant transition to the first excited state only. The S factor obtained in the folding potential calculation of Ref. [7] is used as a reference, S_{ref} . The parameter ξ refers to the strength of the scattering potential.

E (keV)	90	500	1000	
$S_{\text{ref}}(E)$ (keV b)	7.03	4.20	3.39	
$S_{\text{ref}}(E)/S_{\text{ref}}(1 \text{ MeV})$	2.09	1.24	$\equiv 1.00$	
ξ	$S(E)/S_{\text{ref}}(E)$			χ_{red}^2
0.8	1.03	1.07	1.01	1.2
0.9	1.05	1.03	1.01	1.1
1.0	$\equiv 1.00$	$\equiv 1.00$	$\equiv 1.00$	2.2
1.1	0.96	0.97	0.99	4.7

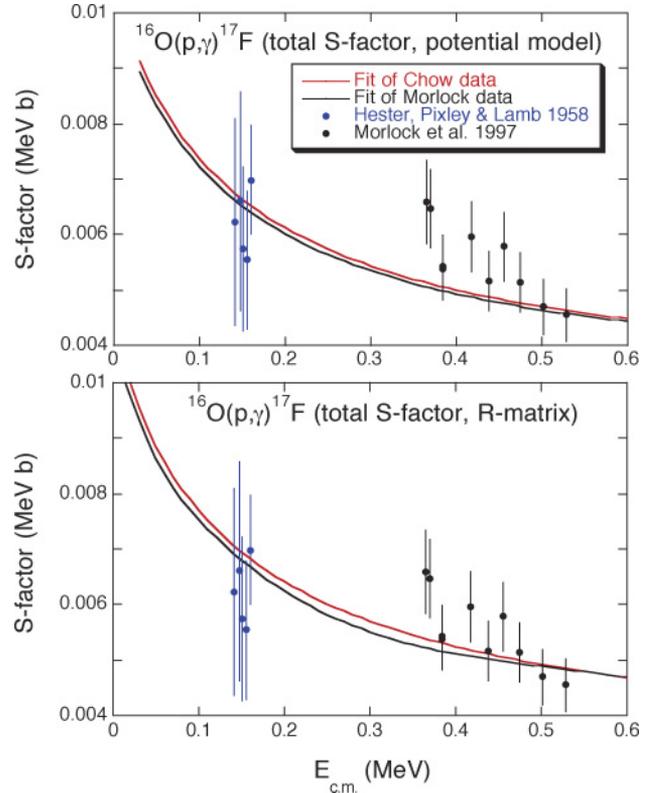


FIG. 5. (Color online) Total S factor for $^{16}\text{O}(p, \gamma)^{17}\text{F}$ at low bombarding energies. The red and black lines indicate fits to the data of Chow, Griffiths, and Hall [6] and Morlock *et al.* [7], respectively, by using a folding potential model (top part; see Sec. III A) and an R -matrix model (bottom part; see Sec. III B). The low-energy data of Hester, Pixley, and Lamb [4] are shown for comparison.

pendence of the S factor, which, in turn, gives rise to a noticeable variation in the χ_{red}^2 value for the adjustment of the spectroscopic factor C^2S . A prerequisite for such a study is a large number of data points as provided by Ref. [7]. For the folding potential we find a broad χ^2 minimum near $\xi \approx 0.9$. The value of χ_{red}^2 increases by 1 if ξ is changed by about 10%. Thus, the energy dependence of the S factor confines the potential strength parameter to the range explored in Table I.

In summary, a variation of the potential parameter ξ within a reasonable physical range (10%) changes the extrapolated S factor by ≈ 1 –5%, depending on the energy E . This value can be compared to the statistical error obtained from the least-squares fits alone (0.5–4%) and to the difference between the absolute normalizations of the two data sets under consideration (3–12%).

The sum of the S factors for both transitions is shown in Fig. 5 (top). The red and black lines indicate the fits to the data of Chow, Griffiths, and Hall [6] and Morlock *et al.* [7], respectively. It can be seen that the analyses of both data sets give consistent results. As a cross-check, the low-energy data of Hester, Pixley, and Lamb [4] are also shown. Recall that these authors report only the total S factor. Clearly, the solid lines are consistent with the data of Ref. [4], providing further support to our results.

B. R -matrix model

1. Formalism

The R -matrix theory combines the flexibility of a phenomenological approach with the physical content of the Coulomb wave functions. It can be applied to resonant as well as to nonresonant reactions. In the R -matrix theory the relative coordinate R between the colliding nuclei is divided into two sectors, limited by the channel radius a . In the internal region ($r \leq a$) the physics of the problem is determined from the properties of poles. In the external region ($r > a$), the relative wave function is given by a linear combination of Coulomb wave functions. This procedure provides an accurate extrapolation down to low energies, where Coulomb effects are expected to play a major role. The poles are associated with resonances or bound states of the unified nucleus, and their properties (energy, particle, and γ -ray partial widths) are fitted to the available data.

The formalism of the R -matrix method is well known [32]. Here we use the formulation of Ref. [24]. The electromagnetic transition matrix element U^γ involved in Eq. (3) is divided into two parts, associated with the internal and external regions, as

$$U^\gamma = U_{\text{int}}^\gamma + U_{\text{ext}}^\gamma. \quad (6)$$

We present the different contributions in Ref. [24]. The internal part U_{int}^γ is given by the energies and partial widths of the poles, whereas the external contribution U_{ext}^γ is provided by an integral involving Coulomb wave functions and is evaluated from the channel radius a to infinity. This external term is proportional to the asymptotic normalization constant (ANC) of the final (bound) state. Let us point out that both contributions are related to each other, as the external term involves the phase shift, which, in turn, depends on the R matrix.

The R -matrix theory can be easily extended to nonresonant reactions. In that case one usually assumes that the nonresonant contribution is simulated by a high-energy background pole. Although this method is acceptable, it presents some drawbacks: the results are somewhat sensitive to the choice of the pole energy, which in addition has no genuine physical meaning. This problem can be addressed by assuming that the R matrix is constant. Nonresonant transfer reactions have been investigated previously in that way [33]. The R matrix is therefore given by

$$R(E) = R_0^p, \quad (7)$$

where subscript p refers to the proton channel. When $R_0^p = 0$, the problem is equivalent to the hard-sphere approximation.

This procedure can be extended to capture reactions. Let us consider the internal contribution in Eq. (6). According to Ref. [24], we have in a general case

$$U_{\text{int}}^\gamma = i^\ell \exp(i\delta_{\text{HS}}) \frac{1}{1 - LR} \sum_{i=1}^N \frac{\sqrt{\tilde{\Gamma}_i \tilde{\Gamma}_{\gamma,i}}}{E_i - E}, \quad (8)$$

where N is the number of poles, $(E_i, \tilde{\Gamma}_i, \tilde{\Gamma}_{\gamma,i})$ are their energies, particle, and γ -ray partial widths, δ_{HS} is the hard-sphere phase shift, and the constant L is related to the Coulomb

wave functions (see Refs. [24,32] for details). Including the energy dependence of the widths we have

$$\tilde{\Gamma}_i \tilde{\Gamma}_{\gamma,i} = 2P_\ell(E) \tilde{\gamma}_i^2 \tilde{\Gamma}_{\gamma,i}^0 (E - E_f)^{2\lambda+1}, \quad (9)$$

where E_f is the energy of the final state, λ is the order of the multipole, and $\tilde{\gamma}_i^2, \tilde{\Gamma}_{\gamma,i}^0$ are energy-independent quantities. For a nonresonant reaction ($N = 1$), Eq. (8) can be rewritten as

$$U_{\text{int}}^\gamma = i^\ell \exp(i\delta_{\text{HS}}) \frac{1}{1 - LR_0^p} \times [2R_0^p R_0^\gamma P_\ell(E) (E - E_f)^{2\lambda+1}]^{1/2}, \quad (10)$$

where R_0^γ is a constant connected with the electromagnetic transition. This parametrization is equivalent to the usual ‘‘background’’ term but without the redundancy associated with the energy of the background pole. Note that a nonresonant capture transition is characterized by the fitting parameters R_0^p, R_0^γ and the ANC of the final state.

2. Results

For the ground-state transition in $^{16}\text{O}(p, \gamma)^{17}\text{F}$, we have $\ell_i = 1$ ($J_i = 3/2^-$) and a small contribution from $\ell_i = 3$ ($J_i = 5/2^-, 7/2^-$). All these transitions are nonresonant. For the excited-state contribution ($J_f = 1/2^+$), we include the $1/2^-$ resonance at $E_r = 2.50$ MeV, with parameters taken from Ref. [1]. The $J_i = 3/2^-$ partial wave is also included as a nonresonant contribution. For simplicity we assume that the partial waves of the same spin-orbit doublet (that is, $J_i = 5/2^-$ and $J_i = 7/2^-$) have the same R -matrix parameters. For the channel radius we used a value of $a = 5$ fm. A simultaneous fit for the ground and excited states is then performed.

The R -matrix fits for the transitions to the ^{17}F ground and first excited state are shown in Figs. 6 and 7, respectively. For the asymptotic normalization constant of the ground-state transition we find values of $\text{ANC} = 1.19 \pm 0.02 \text{ fm}^{-1/2}$ ($\chi_{\text{red}}^2 = 0.2$) for the data of Chow, Griffiths, and Hall [6] and $\text{ANC} = 1.13 \pm 0.01 \text{ fm}^{-1/2}$ ($\chi_{\text{red}}^2 = 0.6$) for the data of Morlock *et al.* [7]. The absolute normalization of these two data sets differs by 5%. For the transition to the first excited ^{17}F state, the results are $\text{ANC} = 81.0 \pm 0.9 \text{ fm}^{-1/2}$ ($\chi_{\text{red}}^2 = 1.6$) for the data of Ref. [6] and $\text{ANC} = 82.3 \pm 0.3 \text{ fm}^{-1/2}$ ($\chi_{\text{red}}^2 = 0.7$) for the data of Ref. [7], where the absolute normalization differs by 2%. Note that the asymptotic normalization constants from our R -matrix analysis are consistent with the microscopic results of Ref. [34].

The sum of the S factors for both transitions is displayed in Fig. 5 (bottom). The red and black lines show the results from fitting the data of Chow, Griffiths, and Hall [6] and Morlock *et al.* [7], respectively. It is apparent that the R -matrix analyses of both data sets give consistent results and describe the total S -factor data rather well. The channel radius is adopted here as $a = 5$ fm, a typical value used in the literature [24]. The sensitivity of the extrapolated S factor to systematic variations of the channel radius a were investigated at a center-of-mass energy of 90 keV. We find that a variation of the channel radius by 10% changes the extrapolated S factor at $E = 90$ keV by about 2%.

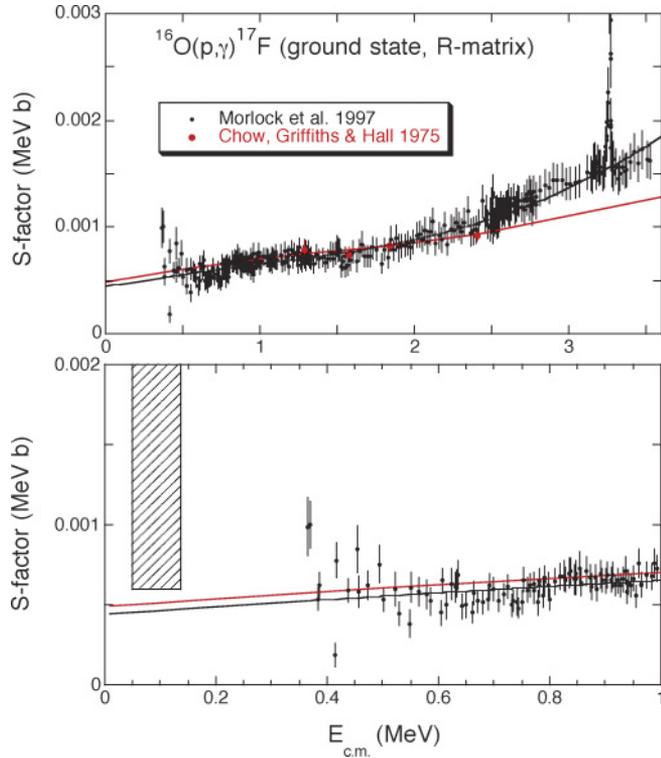


FIG. 6. (Color online) S factor for the proton capture into the ^{17}F ground state. The black and red lines represent fits to the data of Morlock *et al.* [7] and Chow, Griffiths, and Hall [6], respectively. The lines are obtained with an R -matrix model (Sec. III B). The shaded area in the lower part indicates the effective energy range for the hot bottom burning in AGB stars.

IV. RECOMMENDED S -FACTOR AND REACTION RATES

In the previous sections we presented the total S factors and their components for two different data sets, analyzed by using two different nuclear reaction models. We will now describe our method of estimating recommended S factors and their associated errors. This information will then be used to calculate new recommended reaction rates.

The method is explained in Fig. 8 showing, as an example, the total S factor at an energy of $E = 0.090$ MeV, that is, at the center of the effective energy range important for the hot bottom burning in AGB stars. The left and middle panels present results obtained from the folding potential model and the R -matrix approach, respectively. The x-axis labels “CH” and “MO” refer to the analysis of the data of Refs. [6] and [7], respectively. The solid bars indicate 1σ errors resulting from the fitting of the data alone (statistical errors). Note that in general the errors obtained in the R -matrix approach with three fitting parameters in our case are larger than those derived from the potential model with just one fitting parameter. The errors indicated by the open bars are obtained if systematic model uncertainties (i.e., those obtained by varying the parameters of the model within a reasonable physical range) are added quadratically to the statistical errors. The numerical values of the statistical and total errors (in percentages) are listed at the top and bottom, respectively, of each error bar. The label “AV” refers, for each model, to the weighted average of the two data

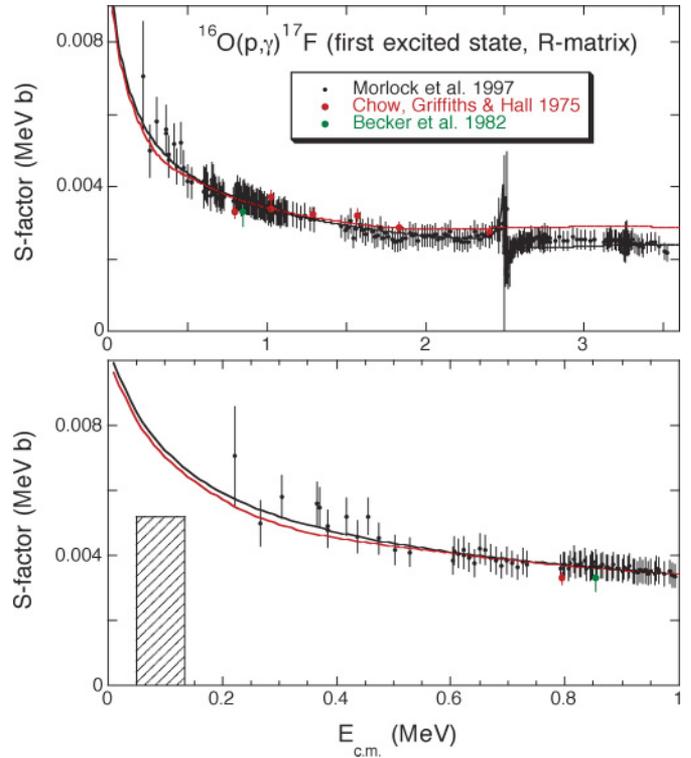


FIG. 7. (Color online) S factor for the proton capture into the first excited state of ^{17}F . The black and red lines represent fits to the data of Morlock *et al.* [7] and Chow, Griffiths, and Hall [6], respectively. The lines are obtained with an R -matrix model (Sec. III B). The shaded area in the lower part indicates the effective energy range for the hot bottom burning in AGB stars.

sets, where only statistical errors are used for the weighing procedure and systematic model errors are added quadratically afterward. The right panel displays the final adopted S factor (label “AD”). Its total error is given by the extrema of the average values resulting from the two reaction models. The mean value of the final recommended S factor is then given by the arithmetic average of the upper and lower bound of the error bar (open bar in right panel).

It was pointed out in Sec. II B that an additional uncertainty of 10% has to be taken into account for the data of Morlock *et al.* Unfortunately, insufficient information is provided in Ref. [7] to determine how much of this uncertainty arises from systematic and from statistical effects. Recall that in the above analysis we treat the entire additional uncertainty as a statistical error in the fitting procedure (see Sec. III). Our treatment will tend to underestimate the true error. Alternatively, one may regard the entire additional uncertainty as a systematic error, which would tend to overestimate the true error. Clearly, the true error must be located somewhere between the two extremes. Test calculations were performed to quantify this effect. Consider as an example the folding potential analysis of the dominating transition to the first excited state for an energy of $E = 0.090$ MeV (similar to the left-hand side of Fig. 8). The alternative assumption of treating the additional uncertainty as a systematic error would increase the averaged (“AV”) S factor by 2.6%, whereas its uncertainty would increase from 5.1 to

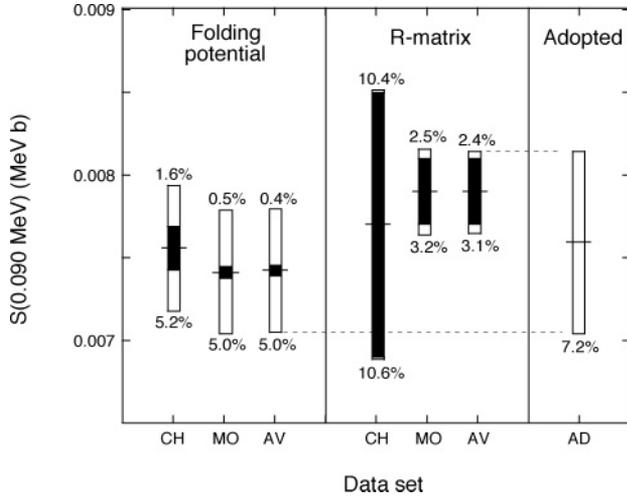


FIG. 8. Total S factor for $^{16}\text{O}(p, \gamma)^{17}\text{F}$ at a center-of-mass bombarding energy of $E = 0.090$ MeV, shown here to illustrate the estimation of the recommended S factor. For an explanation of the x-axis labels and other details, see text.

5.3%. Similar results are obtained in the R -matrix analysis. It is obvious that such small variations are negligible for the adopted S factor (“AD”).

The thermonuclear rate per particle pair for a reaction involving two nuclei is given by [3]

$$N_A \langle \sigma v \rangle = N_A \frac{\sqrt{8/(\pi m)}}{(kT)^{3/2}} \int_0^\infty S(E) e^{-E/kT - 2\pi\eta} dE, \quad (11)$$

where $S(E)$ is the S factor at energy E and the factor $e^{-E/kT}$ derives from the Maxwell-Boltzmann distribution. The quantities T , k , N_A , and m are the plasma temperature, the Boltzmann constant, Avogadro’s constant, and the reduced mass of the interacting nuclei, respectively. The thermonuclear rates for the $^{16}\text{O}(p, \gamma)^{17}\text{F}$ reaction, calculated by integrating

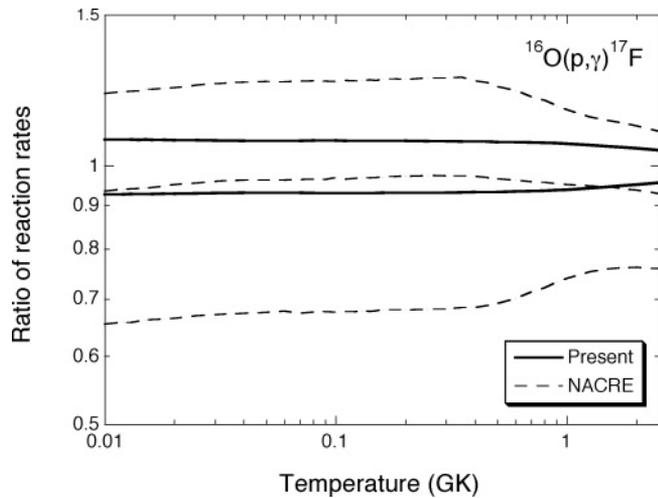


FIG. 9. Reaction rates of $^{16}\text{O}(p, \gamma)^{17}\text{F}$. For better comparison, we show the lower bound, recommended rate and upper bound on the rate normalized to the present recommended rate. Solid and dashed lines indicate the rate ratios for the present and the NACRE [9] results, respectively.

TABLE II. Total thermonuclear reaction rates for $^{16}\text{O}(p, \gamma)^{17}\text{F}$ in units of $(\text{cm}^3 \text{mol}^{-1} \text{s}^{-1})$.

T (GK)	Low	Recommended	High
0.01	6.674×10^{-25}	7.200×10^{-25}	7.733×10^{-25}
0.011	7.026×10^{-24}	7.578×10^{-24}	8.138×10^{-24}
0.012	5.638×10^{-23}	6.080×10^{-23}	6.528×10^{-23}
0.013	3.626×10^{-22}	3.910×10^{-22}	4.197×10^{-22}
0.014	1.941×10^{-21}	2.093×10^{-21}	2.246×10^{-21}
0.015	8.910×10^{-21}	9.604×10^{-21}	1.031×10^{-20}
0.016	3.587×10^{-20}	3.866×10^{-20}	4.148×10^{-20}
0.018	4.211×10^{-19}	4.536×10^{-19}	4.866×10^{-19}
0.02	3.505×10^{-18}	3.775×10^{-18}	4.048×10^{-18}
0.025	2.431×10^{-16}	2.616×10^{-16}	2.803×10^{-16}
0.03	6.124×10^{-15}	6.586×10^{-15}	7.053×10^{-15}
0.04	6.659×10^{-13}	7.155×10^{-13}	7.657×10^{-13}
0.05	1.847×10^{-11}	1.984×10^{-11}	2.123×10^{-11}
0.06	2.309×10^{-10}	2.481×10^{-10}	2.655×10^{-10}
0.07	1.726×10^{-9}	1.855×10^{-9}	1.985×10^{-9}
0.08	9.029×10^{-9}	9.706×10^{-9}	1.039×10^{-8}
0.09	3.644×10^{-8}	3.917×10^{-8}	4.193×10^{-8}
0.1	1.208×10^{-7}	1.299×10^{-7}	1.391×10^{-7}
0.11	3.439×10^{-7}	3.697×10^{-7}	3.957×10^{-7}
0.12	8.662×10^{-7}	9.312×10^{-7}	9.966×10^{-7}
0.13	1.975×10^{-6}	2.124×10^{-6}	2.273×10^{-6}
0.14	4.149×10^{-6}	4.460×10^{-6}	4.773×10^{-6}
0.15	8.133×10^{-6}	8.742×10^{-6}	9.355×10^{-6}
0.16	1.504×10^{-5}	1.616×10^{-5}	1.729×10^{-5}
0.18	4.449×10^{-5}	4.781×10^{-5}	5.114×10^{-5}
0.2	1.129×10^{-4}	1.213×10^{-4}	1.298×10^{-4}
0.25	7.231×10^{-4}	7.766×10^{-4}	8.305×10^{-4}
0.3	2.954×10^{-3}	3.172×10^{-3}	3.391×10^{-3}
0.35	9.035×10^{-3}	9.696×10^{-3}	1.036×10^{-2}
0.4	2.262×10^{-2}	2.426×10^{-2}	2.592×10^{-2}
0.45	4.896×10^{-2}	5.250×10^{-2}	5.605×10^{-2}
0.5	9.495×10^{-2}	1.018×10^{-1}	1.086×10^{-1}
0.6	2.813×10^{-1}	3.012×10^{-1}	3.212×10^{-1}
0.7	6.659×10^{-1}	7.122×10^{-1}	7.587×10^{-1}
0.8	$1.350 \times 10^{+0}$	$1.442 \times 10^{+0}$	$1.535 \times 10^{+0}$
0.9	$2.447 \times 10^{+0}$	$2.610 \times 10^{+0}$	$2.775 \times 10^{+0}$
1.0	$4.074 \times 10^{+0}$	$4.340 \times 10^{+0}$	$4.607 \times 10^{+0}$
1.25	$1.123 \times 10^{+1}$	$1.192 \times 10^{+1}$	$1.261 \times 10^{+1}$
1.5	$2.418 \times 10^{+1}$	$2.557 \times 10^{+1}$	$2.696 \times 10^{+1}$
1.75	$4.437 \times 10^{+1}$	$4.677 \times 10^{+1}$	$4.918 \times 10^{+1}$
2.0	$7.295 \times 10^{+1}$	$7.667 \times 10^{+1}$	$8.041 \times 10^{+1}$
2.5	$1.586 \times 10^{+2}$	$1.659 \times 10^{+2}$	$1.731 \times 10^{+2}$

Eq. (11) numerically using the S factor recommended in this work, are listed in Table II for temperatures in the range of $T = 0.01$ – 2.5 GK. For higher temperatures, the reaction rates are influenced by energies not covered in the present work ($E > 2.4$ MeV).

Results for the reaction rates of $^{16}\text{O}(p, \gamma)^{17}\text{F}$ are shown in Fig. 9. The curves display the ratios of lower bound, recommended rate, and upper bound on the rate over the present recommended rate. Solid and dashed lines indicate the rate ratios resulting from our work and from NACRE [9], respectively. It can be seen that the ratio of the NACRE and the present *recommended* rates is close to unity. However, below

a temperature of $T = 0.5$ GK, the present reaction rate errors amount to only $\approx 7\%$. This represents an improvement by a factor of ≈ 4 over the results published by NACRE, which followed a rather conservative rate evaluation procedure. Note that because of extensive efforts, the $^{16}\text{O}(p, \gamma)^{17}\text{F}$ reaction has now the most precisely known rate involving any target nucleus in the mass $A \geq 12$ range. Astrophysical consequences for the oxygen isotopic ratios from hot bottom burning in AGB stars are presented in the following section.

V. HOT BOTTOM BURNING IN AGB STARS

The aim of this section is to test the effect of our new $^{16}\text{O}(p, \gamma)^{17}\text{F}$ reaction rate on the $^{17}\text{O}/^{16}\text{O}$ ratios produced by HBB in massive AGB stars. This is relevant because of the presence of a very small fraction of stellar meteoritic oxide grains that might have originated in massive AGB stars. These stars may have contributed to the abundances of some radioactive nuclei in the early solar system, for example, ^{26}Al and ^{60}Fe , as well as to other abundance anomalies observed in meteoritic material, for example, ^{62}Ni , ^{87}Rb , and ^{96}Zr [35]. It is important to verify if there is also a population of meteoritic stellar grains that may have originated from this stellar site.

The few (in fact, four) grain candidates for a massive AGB origin exhibit number abundance ratios of $^{18}\text{O}/^{16}\text{O} < 10^{-4}$, $^{17}\text{O}/^{16}\text{O}$ between 7×10^{-4} and 1.5×10^{-3} , and large ^{26}Mg excesses caused by the decay of ^{26}Al , with inferred initial $^{26}\text{Al}/^{27}\text{Al}$ ratios of $\simeq 0.01$ – 0.09 . One of them, spinel grain OC2, also shows an excess in ^{25}Mg . Such signatures could be produced in massive AGB stars via the combined activation of HBB and the operation of the $^{22}\text{Ne} + \alpha$ reaction during unstable He burning (thermal pulses). For a detailed discussion, see Lugaro *et al.* [13]. We focus here on the $^{17}\text{O}/^{16}\text{O}$ ratio, which is the most problematic signature to be interpreted as produced by HBB (see also Ref. [36]).

During HBB the abundance of ^{17}O quickly reaches equilibrium. Thus the $^{17}\text{O}/^{16}\text{O}$ ratio is mostly determined by the ratio of the reaction rates that produce and destroy ^{17}O , that is, the $^{16}\text{O}(p, \gamma)^{17}\text{F}$ and $^{17}\text{O}(p, \alpha)^{14}\text{N}$ reactions, respectively. Because the laboratory measurements of the composition of meteoritic stellar grains provide isotopic ratios of very high precision, with error bars on the order of a few percentages only, small uncertainties in the reaction rates are indispensable

in this case to address the possible match of the models with the observations.

We consider here a stellar model of a typical massive ($6.5 M_{\odot}$) AGB star of solar metallicity ($Z = 0.02$) experiencing HBB. This model was computed by Karakas and Lattanzio [37] using mass loss on the AGB from Ref. [38] and it experienced 39 thermal pulses. The temperature at the base of the convective envelope during the AGB phase increases from $T = 60$ MK to $T = 87$ MK at the 25th thermal pulse and then decreases again to $T = 20$ MK at the end of the computed evolution. The high temperature reached by this model allows for the production of enough ^{26}Al to match the composition of grain OC2, whereas the $^{17}\text{O}/^{16}\text{O}$ ratio could be reproduced within the previous [9] $^{16}\text{O}(p, \gamma)^{17}\text{F}$ reaction rate uncertainties (see Fig. 7 of Ref. [13]).

Similar to the procedure of Lugaro *et al.* [13], we analyze the impact of the reaction rate uncertainties using a post-processing nucleosynthesis code. Hence we disregard any feedback of the different reaction rates on the stellar structure. This is justified in this case because (i) the $^{16}\text{O}(p, \gamma)^{17}\text{F}$ and $^{17}\text{O}(p, \alpha)^{14}\text{N}$ reactions generate only a small amount of energy within the CNO cycle and (ii) the $^{16}\text{O}(p, \gamma)^{17}\text{F}$ reaction is only marginally activated in our model (only 23% of the initial ^{16}O is destroyed by the second dredge-up and by HBB). For the $^{16}\text{O}(p, \gamma)^{17}\text{F}$ reaction we use our new rates, whereas for the $^{17}\text{O}(p, \alpha)^{14}\text{N}$ reaction we employ the same rate as in Ref. [13], which is based on the latest available experimental information [39,40]. Different runs of the $6.5 M_{\odot}$ stellar model are then computed by varying independently the rates of these two reactions between their lower and upper bounds. The results of this procedure are given in Table III, showing how the errors in the rates propagate to uncertainties of the computed $^{17}\text{O}/^{16}\text{O}$ ratio. Note that the calculated $^{17}\text{O}/^{16}\text{O}$ ratio is almost the same as in Ref. [13] because the *recommended* $^{16}\text{O}(p, \gamma)^{17}\text{F}$ reaction rate did not change significantly (see Fig. 9). The main result here is that the total uncertainty in the derived abundance ratio is reduced from +63% and -73% when using the NACRE rate for $^{16}\text{O}(p, \gamma)^{17}\text{F}$ to +35% and -30% when employing our new reaction rate.

In conclusion, the measured $^{17}\text{O}/^{16}\text{O}$ ratio of grain OC2 ($= 1.25 \pm 0.07 \times 10^{-3}$) could be reproduced within the large error bars of the NACRE compilation ($2.44_{-1.78}^{+1.54} \times 10^{-3}$) in models of massive AGB stars; however, the much more precise

TABLE III. Final $^{17}\text{O}/^{16}\text{O}$ abundance ratios for $6.5 M_{\odot}$, $Z = 0.02$ stellar models computed with all possible combinations of $^{16}\text{O}(p, \gamma)^{17}\text{F}$ and $^{17}\text{O}(p, \alpha)^{14}\text{N}$ reaction rates: LR = lower bound of rate (at $T = 60$ – 100 MK), RR = recommended rate, and UR = upper bound on rate (at $T = 60$ – 100 MK). The $^{17}\text{O}/^{16}\text{O}$ ratio is explicitly given for the RR-RR combination in the center of the table, whereas percentages variations with respect to this case are given for the other rate combinations. Results for the $^{16}\text{O}(p, \gamma)^{17}\text{F}$ rate derived in this article are in roman; results using the NACRE [9] rate are in italics.

		$^{16}\text{O}(p, \gamma)^{17}\text{F}$		
		LR (-7% , -43%)	RR	UR ($+7\%$, $+30\%$)
$^{17}\text{O}(p, \alpha)^{14}\text{N}$	LR ($\simeq -22\%$)	$+17\%$, -14%	$+26\%$, $+25\%$	$+35\%$, $+63\%$
	RR	-7% , -43%	2.52×10^{-3} , 2.44×10^{-3}	$+7\%$, $+30\%$
	UR ($\simeq +18\%$)	-30% , -73%	-20% , -21%	-13% , $+7\%$

$^{16}\text{O}(p, \gamma)^{17}\text{F}$ rate of the present work leads to $2.52_{-0.76}^{+0.88} \times 10^{-3}$ for the $^{17}\text{O}/^{16}\text{O}$ ratio and disagrees with the measured value. Consequently, there is not clear evidence to date for any stellar grain origin from massive AGB stars. Stellar model uncertainties, such as different mixing prescriptions [41] and mass loss rates [42] still need to be carefully evaluated in this context. Another possibility is that we have not yet discovered grains from massive AGB stars because they are perhaps smaller in size than the grains currently analyzed in the laboratory ($>1 \mu\text{m}$) [43,44]. Clearly, a solution to this problem requires future efforts.

VI. SUMMARY

In this work we focused our attention on a reanalysis of the $^{16}\text{O}(p, \gamma)^{17}\text{F}$ reaction rates. We started from an evaluation of all the original data and performed a number of corrections when appropriate. The modified and improved data are then interpreted in terms of two independent models of nuclear reactions, a potential model and an R -matrix approach. This attention to detail was clearly not practical in previous reaction

rate evaluations. After combining the results from the two reaction models, we find $^{16}\text{O}(p, \gamma)^{17}\text{F}$ reaction rate errors of less than 7% over the entire range of astrophysical interest. In other words, the $^{16}\text{O}(p, \gamma)^{17}\text{F}$ reaction exhibits now the most precisely determined thermonuclear rates among any charged-particle capture reactions in the $A \geq 12$ mass range. Our new results are then incorporated into models of massive AGB stars to study the derived oxygen isotopic ratios. Contrary to previous conclusions, we find now that there is presently no clear evidence of a massive AGB star origin for any observed stellar grains.

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